

ON THE OCCURRENCE OF WR+O BINARIES.

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Using stellar evolution models of two different kinds, one with convective cores determined by the Schwarzschild criterion, one with enlarged convective cores, we discuss the **masses of the WR stars** in a number of double-lined WR binaries. For the main group of systems the mass of the WR-star is between 10 and 20 M_{\odot} .

We further show that the virtual **absence of WR+B systems** can be explained as a result of accretion effects on the inner structure of the companion star (rejuvenation) and of the short timescale of existence of the WR star relative to the dominant timescale of the companion star. Systems with WR components that were transformed directly by stellar wind may now possibly be observed as WR-SB1 systems with a large mass function.

Examination of the formation mechanism of WR stars in binaries (stellar wind or mass transfer) leads to the conclusion that **mass transfer is probably the acting mechanism** in most of the presently **well-known systems**. Especially this must be the case for the WR binaries in the Small Magellanic Cloud.

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